



Politecnico
di Bari

Repository Istituzionale dei Prodotti della Ricerca del Politecnico di Bari

Detection of a thermal spectral component in the prompt emission of GRB 100724B

This is a pre-print of the following article

Original Citation:

Detection of a thermal spectral component in the prompt emission of GRB 100724B / Guiriec, Sylvain; Connaughton, Valerie; Briggs, Michael S.; Burgess, Michael; Ryde, Felix; Daigne, Frédéric; Mészáros, Peter; Goldstein, Adam; Mcenery, Julie; Omodei, Nicola; Bhat, P. N.; Bissaldi, Elisabetta; Camero-Arranz, Ascensión; Chaplin, Vandiver; Diehl, Roland; Fishman, Gerald; Foley, Suzanne; Gibby, Melissa; Giles, Misty M.; Greiner, Jochen; Gruber, David; von Kienlin, Andreas; Kippen, Marc; Kouveliotou, Chryssa; Mcbreen, Sheila; Meegan, Charles A.; Paciesas, William; Preece, Robert; Rau, Arne; Tierney, Dave; van der Horst, Alexander J.; Wilson-Hodge, Colleen. - In: THE ASTROPHYSICAL JOURNAL LETTERS. - ISSN 2041-8205. - STAMPA. - 727:2(2011). [10.1088/2041-8205/727/2/L33]
Availability: This version is available at <http://hdl.handle.net/11589/205168> since: 2021-03-09

Published version

DOI:10.1088/2041-8205/727/2/L33

Publisher:

Terms of use:

(Article begins on next page)

DETECTION OF A THERMAL SPECTRAL COMPONENT IN THE PROMPT EMISSION OF GRB 100724B

SYLVAIN GUIRIEC¹, VALERIE CONNAUGHTON¹, MICHAEL S. BRIGGS¹, MICHAEL BURGESS¹, FELIX RYDE^{2,3}, FRÉDÉRIC DAIGNE^{4,17},
PETER MÉSZÁROS⁵, ADAM GOLDSTEIN¹, JULIE MCENERY^{6,7}, NICOLA OMODEI⁸, P.N. BHAT¹, ELISABETTA BISSALDI⁹,
ASCENSIÓN CAMERO-ARRANZ¹⁰, VANDIVER CHAPLIN¹, ROLAND DIEHL⁹, GERALD FISHMAN¹¹, SUZANNE FOLEY⁹,
MELISSA GIBBY¹², MISTY M. GILES¹², JOCHEN GREINER⁹, DAVID GRUBER⁹, ANDREAS VON KIENLIN⁹, MARC KIPPEN¹³,
CHRYSSA KOUVELIOTOU¹¹, SHEILA MCBREEN¹⁴, CHARLES A. MEEGAN¹⁵, WILLIAM PACIESAS¹, ROBERT PREECE¹, ARNE RAU⁹,
DAVE TIERNEY¹⁴, ALEXANDER J. VAN DER HORST^{16,18}, AND COLLEEN WILSON-HODGE¹¹

¹ University of Alabama in Huntsville, NSSTC, 320 Sparkman Drive, Huntsville, AL 35805, USA; sylvain.guiriec@nasa.gov, sylvain.guiriec@uah.edu

² Department of Physics, Royal Institute of Technology, AlbaNova, SE-106 91 Stockholm, Sweden

³ The Oskar Klein Centre for Cosmo Particle Physics, AlbaNova, SE-106 91 Stockholm, Sweden

⁴ Institut d'Astrophysique de Paris, UMR 7095, Université Pierre et Marie Curie-Paris 06, CNRS 98 bis boulevard Arago, 75014 Paris, France

⁵ Department of Astronomy & Astrophysics, Department of Physics and Center for Particle Astrophysics, Pennsylvania State University, University Park, PA 16802, USA

⁶ NASA Goddard Space Flight Center, Greenbelt, MD 20771, USA

⁷ Department of Physics and Department of Astronomy, University of Maryland, College Park, MD 20742, USA

⁸ Istituto Nazionale di Fisica Nucleare, Sezione di Pisa, I-56127 Pisa, Italy

⁹ Max-Planck-Institut für extraterrestrische Physik, Giessenbachstrasse 1, 85748 Garching, Germany

¹⁰ National Space Science and Technology Center, 320 Sparkman Drive, Huntsville, AL 35805, USA

¹¹ Space Science Office, VP62, NASA/Marshall Space Flight Center, Huntsville, AL 35812, USA

¹² Jacobs Technology, Inc., Huntsville, AL, USA

¹³ Los Alamos National Laboratory, P.O. Box 1663, Los Alamos, NM 87545, USA

¹⁴ University College, Dublin, Belfield, Stillorgan Road, Dublin 4, Ireland

¹⁵ Universities Space Research Association, NSSTC, 320 Sparkman Drive, Huntsville, AL 35805, USA

¹⁶ NASA/Marshall Space Flight Center, 320 Sparkman Drive, Huntsville, AL 35805, USA

Received 2010 October 20; accepted 2010 November 23; published 2011 January 10

ABSTRACT

Observations of GRB 100724B with the *Fermi* Gamma-Ray Burst Monitor find that the spectrum is dominated by the typical Band functional form, which is usually taken to represent a non-thermal emission component, but also includes a statistically highly significant thermal spectral contribution. The simultaneous observation of the thermal and non-thermal components allows us to confidently identify the two emission components. The fact that these seem to vary independently favors the idea that the thermal component is of photospheric origin while the dominant non-thermal emission occurs at larger radii. Our results imply either a very high efficiency for the non-thermal process or a very small size of the region at the base of the flow, both quite challenging for the standard fireball model. These problems are resolved if the jet is initially highly magnetized and has a substantial Poynting flux.

Key words: acceleration of particles – gamma-ray burst: individual (GRB 100724B) – gamma rays: stars – radiation mechanisms: non-thermal – radiation mechanisms: thermal

1. INTRODUCTION

The prompt emission detected from gamma-ray bursts (GRBs) is believed to originate at large distances from the central engine, from within an ultrarelativistic outflow (Piran 2004). This ultrarelativistic motion is necessary to avoid strong $\gamma\gamma$ annihilation, a signature that is not observed (see, e.g., Piran 1999). Thermal emission is naturally expected in such a scenario. Indeed, since the densities at the base of the relativistic flow are very large, the medium is optically thick to radiation owing to Thomson scattering by entrained electrons. The optical depth decreases during the relativistic expansion and the outflow eventually becomes transparent for its own radiation, at the photospheric radius. Any internal energy that is still carried out by the flow can be radiated at the photosphere and will be observed as a thermal component in the prompt spectrum. This expected photospheric emission in GRB spectra was early suggested on such theoretical grounds by Goodman (1986), Mészáros (2002), and Rees & Mészáros (2005), among others. The non-thermal component observed in the spectrum has to be produced by another mechanism in the optically thin region, i.e., well above

the photosphere. Due to the ultrarelativistic motion, this difference in the radius of the emission implies a delay between the observation of the two components that is usually small compared to the typical duration of a long GRB and is also small compared to the typical duration of time intervals used for time-dependent spectroscopic analysis. The thermal and non-thermal components should then appear superimposed for the observer (e.g., Mészáros & Rees 2000). Daigne & Mochkovitch (2002) pointed out that in the standard fireball model, the photospheric component can easily be dominant in the spectrum if the efficiency f_{NT} of the mechanism responsible for the non-thermal emission is only moderate ($f_{\text{NT}} \lesssim 40\%$).

Observationally, Ghirlanda et al. (2003), Ryde (2004, 2005), and Ryde et al. (2010) argued that a photospheric component is present in CGRO BATSE data. The limited energy range provided by BATSE (20–2000 keV), however, hampered the possibility of unambiguously identifying the emission process. Since the launch of *Fermi* in 2008, the combination of the Gamma-Ray Burst Monitor (GBM) and the Large Area Telescope (LAT) provides an unprecedented energy range for GRB spectroscopy, and the identification of the emission processes responsible for the gamma-ray prompt emission may become a reality. GBM alone covers a wider energy range than its predecessor BATSE,

¹⁷ Institut Universitaire de France

¹⁸ NASA Postdoctoral Program Fellow.

and the design of its data enables finer resolution spectroscopy. This allows better constraints on spectral fits, with increasingly complex models.

GRB energy spectra in the keV–MeV energy range are usually well represented by the Band function (Band et al. 1993), two power laws, smoothly joined, and parameterized by E_{peak} , which represents the energy at which peak power is radiated (Gehrels 1997). The value of the low-energy power-law index, α , is higher than the value of the high-energy power-law index, β , and the parameter E_{peak} of the Band function for GRBs generally appears to follow predictable trends with time and flux level (Ford et al. 1995; Guiriec et al. 2010). It is, however, an empirical function rather than a physically motivated model. The meaning of the parameters in the context of emission and transport mechanisms taking place in GRBs is not well understood, but is generally believed to represent the non-thermal emission from accelerated charged particles.

In Section 1, we describe our observational results consisting of a GRB prompt-emission spectrum best fit with the combination of a thermal component and a standard Band function. In Section 2, we use these results to constrain the origin of the energy released in the GRB jet.

2. OBSERVATIONS

GBM is composed of 12 sodium iodide (NaI) detectors covering an energy range from 8 keV to 1 MeV and two bismuth germanate (BGO) detectors sensitive between 200 keV and 40 MeV (Meegan et al. 2009). The instrument triggered on 2010 July 24, at $T_0 = 00:42:05.992$ UT on the very bright GRB 100724B (Bhat 2010). The event was also seen at higher energies in the *Fermi* Large Area Telescope (LAT; Tanaka et al. 2010). The most precise position for the direction of the burst is the intersection of the InterPlanetary Network annulus obtained using GBM, *Konus-WIND* (Golenetskii et al. 2010), and *MESSENGER* data with the 90% LAT confidence level location error box, and is a strip of sky centered on R.A. = $118^{\circ}8$ and decl. = $75^{\circ}8$ which is $1^{\circ}2$ long and $0^{\circ}2$ wide (K. Hurley & V. Pal'shin 2010, private communication). Figure 1 (top two panels) shows the GBM light curve of GRB 100724B in two energy bands. Multiple peaks of varying intensity are superimposed on a pre-trigger plateau, with a decaying tail that is detected over 200 s from T_0 .

We simultaneously fit the spectral data of the NaI detectors with a source angle less than 60° (NaIs 0, 1, 2, 3, and 5) and the data of the brightest BGO detector (BGO 0) using the analysis package Rmfit 3.3rc8. An effective area correction is applied between each of the NaIs and BGO 0 during the fit process. This correction is used to handle possible discrepancies between the flux in the detectors due to the choice of the model to generate the instrument responses for instance.

We performed a time-integrated spectral analysis over the main part of the burst ($T_0 - 1.024$ s to $T_0 + 83.969$ s) using the Band function. The Band parameters are in part fairly typical of the ensemble of GRBs, with $\alpha = -0.67 \pm 0.01$ and $E_{\text{peak}} = 352 \pm 6$ keV (Preece et al. 2000; Kaneko et al. 2006). However, with an index $\beta = -1.99 \pm 0.02$, the high-energy power law systematically overshoots the observed flux above 1 MeV in BGO, as can be seen by the fit residuals in Figure 2 (top two panels), which also indicate systematic patterns at low energy. This suggests that a simple Band function does not adequately represent the spectrum of this burst.

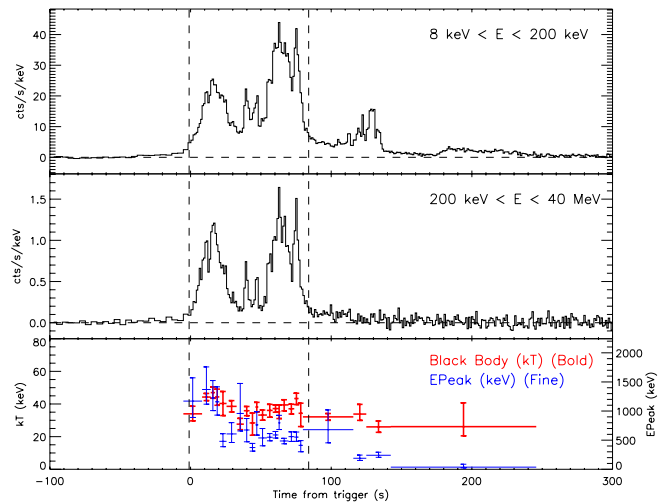


Figure 1. Top two panels show the signal count rates as a function of time, as measured by the *Fermi* GBM detectors, from 8 to 200 keV in NaI (top) and from 200 keV to 40 MeV in BGO (middle). The bottom panel shows the evolution of the Band function E_{peak} (in blue) and the BB temperature kT (in red) over the duration of the burst. The vertical dashed lines indicate the period used in the time-integrated analysis.

We identify the best shape to fit the above-mentioned spectral deviation by fitting the same data simultaneously with a Band function combined with each of the following models: single power law (PL), blackbody (BB), Band function, power law with exponential cut off (“Comp” for Comptonized model), and Gaussian. We select the best model by choosing the fit with the lowest Castor C-stat value (later C-stat). C-stat differs from the Poisson likelihood statistic by an offset which is a constant for a particular data set. Table 1 shows the results of these fits. The effective area correction described above is on the order of a few percent and does not change the C-stat for each fit more than a few units, nor does it change the value of the parameters resulting from the fit.

While spectral deviations from the standard Band function were previously identified in the form of an additional PL to the Band function sometimes extending from the lower energy in the GBM to the higher energy in the LAT (Abdo et al. 2009; Ackermann et al. 2010a, 2010b; Guiriec et al. 2010), in the case of GRB 100724B a PL spectral component does not improve on the Band-only fit and an additional BB component to the Band function is the best model to fit the spectral deviation. An equal C-stat is obtained for Band+Band and Band+Compt, but with α close to +1 for the additional Band and Compt functions, and a very low value for β (only constrained as an upper limit, below -5) for the extra Band function, the Band and Compt functions can be interpreted as a Planck function. Even with more parameters, the additional Band and Compt functions resemble a BB component, reinforcing Band+BB as the best combination. Additional models were tried, such as a log-Parabola function (Massaro et al. 2010), but the results were highly disfavored, and we exclude them from Table 1.

Figure 2 (bottom two panels) shows the BB contribution below E_{peak} . Compared to the Band-only fit, E_{peak} is shifted toward higher energy to 615 ± 29 keV and β is lower with a value of -2.11 ± 0.02 . This index is consistent with the flux detected above 1 MeV, and the spectrum seen in the LAT (Tanaka et al. 2010) at higher energies. α is also significantly lowered to -0.90 ± 0.02 .

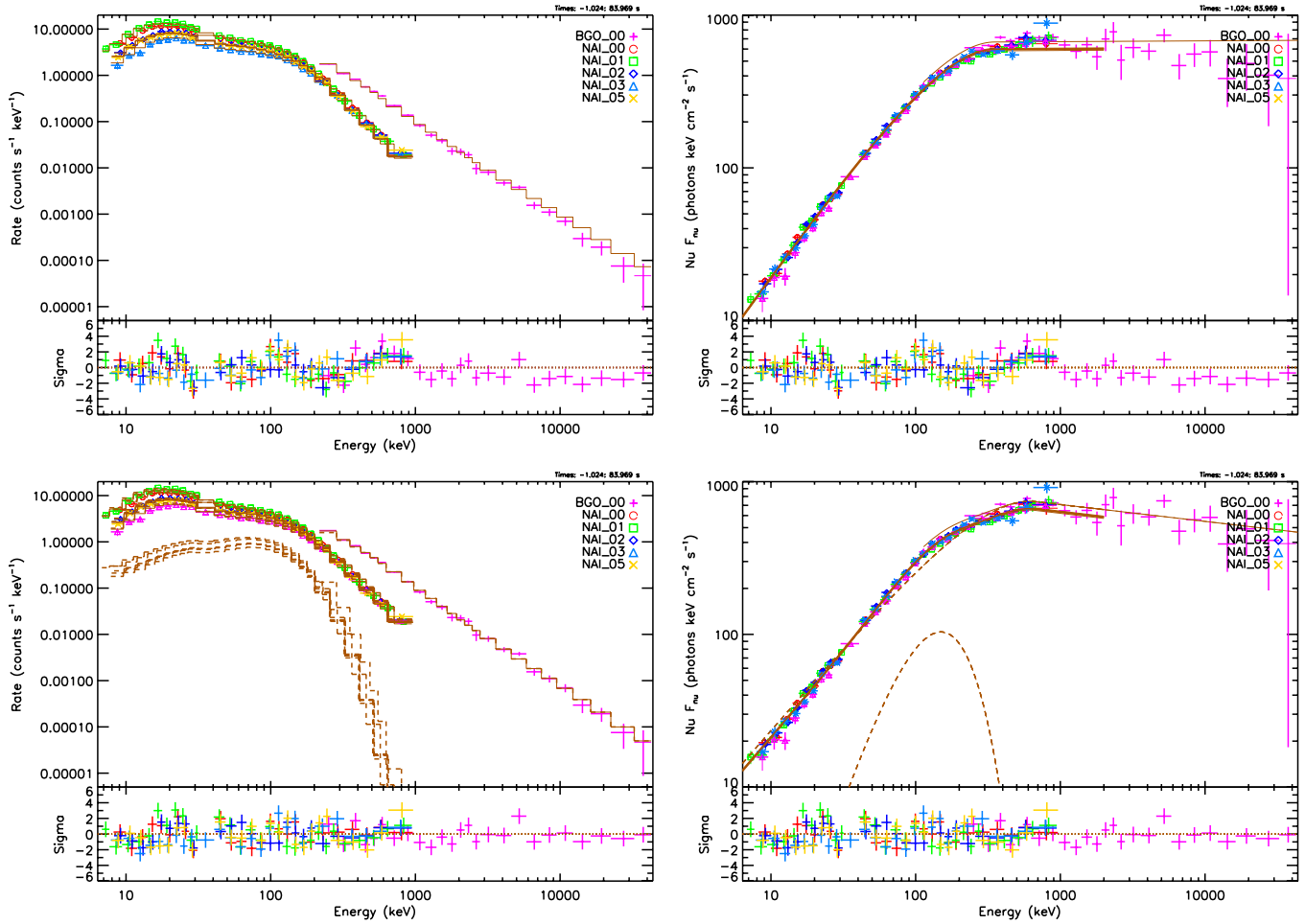


Figure 2. Time-integrated spectrum of GRB 100724B fit by a Band function (top two panels) and a Band+BB function (bottom two panels). The left plots show the count spectra for the two models, and the right plots show the corresponding model fit deconvolved νF_ν spectra. The data points appear as color crosses. Dashed lines indicate the individual spectral functions, and solid lines show the summed model fit. The addition of a BB spectral component over the brightest part of the burst ($T_0-1.024$ s to $T_0+83.969$ s) shows a significant improvement in the fit compared to a Band function by itself, particularly noticeable as the removal of trends with energy in the residuals compared to the Band-only fit. The region between 30 and 40 keV is excluded from the fit owing to calibration issues around the k -edge of the NaI detectors. We have verified that this exclusion does not affect the recovered parameter values.

Table 1
Fit of the Time-integrated Spectrum of GRB 100724B from $T_0-1.024$ s to $T_0+83.969$ s

Models	Standard Model			Additional Model									
	Band		BB	Compt		Band		Gaussian		PL	C-stat/dof		
Parameters	E_{peak}	α	β	kT	E_{peak}	Index	E_0	α	β	Centroid	\log_{10} FWHM	Index	
Band	352	-0.67	-1.99									1133/704	
	± 6	± 0.01	± 0.01										
Band+BB	615	-0.90	-2.11	38.14								1038/702	
	± 29	± 0.02	± 0.02	± 0.87									
Band+Compt	708	-0.94	-2.13		164	+0.81						1039/701	
	± 48	± 0.02	± 0.02		± 7	± 0.20							
Band+Band	716	-0.94	-2.13				60	0.76	< -5			1039/700	
	± 48	± 0.02	± 0.02				± 7	± 0.21					
Band+Gaussian	403	-0.75	-2.02							103	0.25	1060/701	
	± 8	± 0.01	± 0.01							± 2	± 0.03		
Band+PL	341	-0.63	-1.99									-1.93	1131/702
	± 9	± 0.05	± 0.01									± 1.59	

Notes. The count spectrum using the NaI detectors 0, 1, 2, 3, 5 and BGO detector 0 is fit simultaneously with a standard Band function and with an additional model to evaluate the shape of the spectral deviation. Band+BB is preferred over all the other combinations.

While the simultaneous fit of all the selected detectors provides the best constraints on the two spectral components, fits with Band+BB to combinations of individual NaI detectors with

BGO 0 result in similar parameter values and offer significant improvement over the Band-only fit. This provides a check that the BB component is real and not introduced by effects such

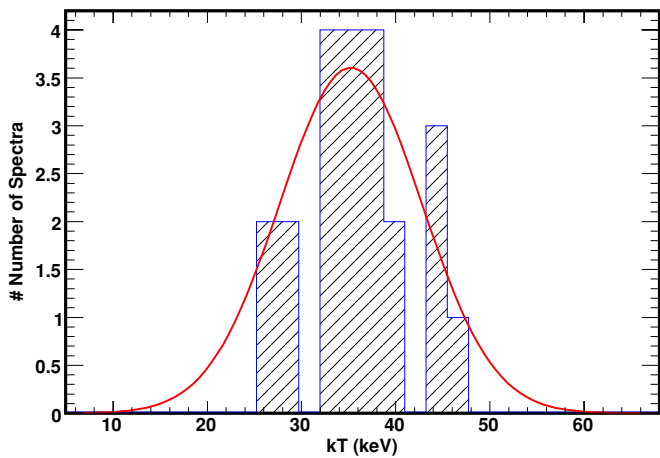


Figure 3. Distribution of the time-resolved BB temperature kT . The spread in temperatures measured in the BB component over the 22 time intervals as shown in Figure 1 can be fit by a Gaussian distribution with a mean of 35.2 ± 2.3 , and a 1σ standard deviation of 6.0 ± 2.0 .

as detector dead time or spectral distortions that would affect each detector in a different way depending on the angle of the detector to the source.

To verify that the improvement in the fit obtained by adding a BB component to the Band function is not a statistical fluctuation, we generated 20,000 synthetic spectra for each selected detector. For the simulations we used the parameters from the fit performed with the Band-only function, which we take as the null hypothesis. To create the simulated spectra, for each detector the real background is added to the source spectrum model and Poissonian fluctuations are applied to the sum. All the detectors are then fit simultaneously with both Band and Band+BB, and their C-stat are compared. None of the 20,000 simulated spectra give a difference larger than 45 units of C-stat ($[\text{Band}] - [\text{Band} + \text{BB}]$), while in the real data, this difference is 95 units, corresponding to a probability lower than 5×10^{-5} that the BB excess is due to statistical fluctuations.

To study the evolution of the spectral components, 22 time intervals were devised by requiring that each interval produces a Band+BB spectral fit with well-constrained Band function parameters, while attempting to separate the peaks and valleys of the light curve so that the spectral fit parameters can be tracked with burst flux as well as with time. The bottom panel of Figure 1 exhibits the evolution of E_{peak} and BB temperature, kT , through these intervals. We notice that E_{peak} tracks the light curve and globally decreases over time. The BB component is detected throughout the burst, and its temperature shows weak correlation with E_{peak} . The significance of this correlation is difficult to assess, mostly because the variation in temperature is small. Overall, it appears that the temperature is quite stable, with Figure 3 showing more clearly the small scatter in kT .

3. DISCUSSION

With a fluence of $\sim 5.2 \times 10^{-4}$ erg cm $^{-2}$ measured in 85 s from $T_0 - 1.024$ s between 8 keV and 40 MeV, GRB 100724B is the most intense GRB detected by GBM over this energy range through 2010 September. Combined with the broad energy range of the GBM, this allows for accurate modeling of its energy spectrum even with this complex model. Previous observational results regarding thermal components in GRBs were ambiguous and some were limited to individual fine time slices rather than a spectral fit over the entire emission period.

Some studies showing BB fits did not demonstrate that the BB fit was statistically preferred to a simple non-thermal component (Ghirlanda et al. 2003). Other analyses found BB+PL spectra for isolated portions of selected GRBs, raising the possibility that these spectra are actually adequately fit with a standard Band function but that due to a weak signal in small time slices and extreme parameters for the Band function, a BB shape is competitive with the Band function (Ryde 2004, 2005; Ryde et al. 2010). The non-thermal component fit with a single power law suggested a break well beyond the common E_{peak} values, and the BB temperature and its variations intriguingly matched those of a typical E_{peak} . Despite the broader energy range of *RHESSI* GRB observations, one analysis found difficulties in fitting combined thermal plus non-thermal models (Bellm 2010).

We find here that the joint BB plus non-thermal (Band) fit is highly statistically preferred so that in simultaneously detecting both components we are confident in their correct identification. Time-resolved spectroscopy of GRB 100724B reveals that this BB component is seen throughout the burst and does not evolve much over time, while the non-thermal component follows the typical variations (Ford et al. 1995; Guiriec et al. 2010). The consistency of the mean kT value with the temperature obtained in the time-integrated spectral fit, combined with the detection of the BB component throughout the burst, strengthen the case for an underlying thermal component in the gamma-ray emission seen from GRB 100724B and show that the presence of the BB in the time-integrated spectrum cannot be attributed to spectral evolution of the Band function during the burst.

E_{peak} varies substantially, from ~ 90 to ~ 1100 keV. At the same time, the thermal component remains relatively steady with the temperature varying only modestly between 30 and 50 keV as suggested by a $T \propto L^{1/2}$ dependence expected from a BB component. Time-averaged values of the temperature and the flux of the thermal component, and of the ratio of this flux over the total gamma-ray flux are $kT = 38 \pm 4$ keV, $F_{\text{bb}} = (2.6 \pm 1.4) \times 10^{-7}$ erg s $^{-1}$ cm $^{-2}$, and $F_{\text{bb}}/F_{\text{tot}} = 0.04 \pm 0.02$.

In the standard fireball model these observables allow determination of the physical properties of the outflow and its photosphere. Owing to the imprecise and delayed localization of GRB 100724B, optical follow up to determine the distance to the source was impossible. For this reason the temperature of the BB can be translated into a real source temperature only as a function of source distance. We assume in the following argument a typical redshift $z = 1$. We find that the Lorentz factor is $\Gamma \simeq 325 \xi^{1/4} f_{\text{NT}}^{-1/4}$, the photospheric radius is $R_{\text{ph}} \simeq 5.6 \times 10^{11}$ cm $\xi^{-3/4} f_{\text{NT}}^{-1/4}$, and the radius at the base of the flow is $R_0 \simeq 1.2 \times 10^7$ cm $\xi^{-1} f_{\text{NT}}^{3/2}$ (Daigne & Mochkovitch 2002; Pe'er et al. 2007). Here, ξ is a geometrical factor of order unity and f_{NT} is the efficiency of the mechanism responsible for the non-thermal emission. With an extreme efficiency $f_{\text{NT}} = 1$, these estimates are in good agreement with the typical values expected in the fireball model. The dependence on redshift is not strong: at $z = 3$ (resp. 8), $\Gamma \simeq 645$ (resp. 1290), $R_{\text{ph}} \simeq 1.1 \times 10^{12}$ cm (resp. 1.4×10^{12} cm), and $R_0 \simeq 1.1 \times 10^7$ cm (resp. 6.9×10^6 cm).

Using more realistic values for the efficiency, the radius R_0 is the most altered, with $R_0 \simeq (3.6\text{--}40 \text{ km}) \xi^{-1}$ for $f_{\text{NT}} \simeq 0.1\text{--}0.5$. Such small values are puzzling. If the central engine is a rotating black hole, as in the popular collapsar model for long GRBs (Woosley 1993), with a minimal mass in the range 5–10 M_{\odot} , such radii are smaller than the typical value expected for the innermost stable orbit, from 44–89 km for a non-rotating black

hole to 22–43 km for a highly rotating black hole having a spin $a = 0.8$. These results for the time-integrated spectrum imply a small R_0 or a very large efficiency and the constraint is even stronger in some time bins. We conclude that observations of GRB 100724B require either a very high efficiency for the non-thermal process or a very small size of the region at the base of the flow, both of which are quite challenging for the standard fireball model, if not excluding it.

A simple solution to this discrepancy between the standard fireball model and the observations is to assume that the initial energy release by the central engine is not purely thermal, but that the flow is highly magnetized close to the source (Daigne & Mochkovitch 2002; Zhang & Pe'er 2009). The magnetization σ is the ratio of the Poynting flux over the power (thermal + kinetic) carried by the baryons. If no magnetic dissipation occurs below the photosphere, the efficiency f_{NT} in the estimates of Γ , R_{ph} , and R_0 above should be replaced by $(1+\sigma)f_{\text{NT}}$. A magnetization $\sigma > 1$ will therefore reconcile the observations with physically acceptable values for the radius at the base of the flow and the efficiency of the mechanism responsible for the non-thermal emission. A similar conclusion is reached for scenarios where magnetic dissipation occurs early and contributes efficiently to the acceleration of the jet. However, the appearance of a low intensity thermal component in the spectrum probably excludes the most extreme version of the magnetized outflow scenario, where the energy is released by the central engine as a pure Poynting flux ($\sigma = \infty$).

4. CONCLUSION

We have shown that the simultaneous presence of thermal and non-thermal components to the spectra of GRB 100724B is statistically preferred. Although the non-thermal component is dominant, the BB flux is well within the GBM sensitivity. Deviations from the Band function may be measurable in less fluent bursts or in bursts where the thermal component is less prominent, providing that the BB component lies in the band pass of the instrument and its peak in energy is distinguishable from E_{peak} . If the presence of an unresolved thermal component in other bursts modifies the Band function parameters in the same sense as the Band-only fit for GRB 100724B, then we might expect a systematic bias yielding values of α and β that are higher (harder) than in the true non-thermal component. Two important consequences of this bias are that the perceived violation of the synchrotron limit that disallows values $\alpha > -2/3$ (for slow-cooling electrons) and $\alpha > -3/2$ (for fast-cooling electrons) may not be as common as suggested by Preece et al. (1998) and Crider et al. (1997), and that the relatively low rate of bursts detected by the LAT compared to the predictions of Band et al. (2009) and the observations in A. A. Abdo et al. (2011, in preparation) based on extrapolations of β from lower energies might be explained by this bias in β , a possibility suggested also by Ryde & Pe'er (2009).

Our observations provide strong evidence for the presence of a photospheric spectral component, long suspected to exist in the standard fireball model. In addition, our results require implausible parameters for the standard baryonic fireball model and therefore favor a substantial magnetic component to the outflow.

The GBM project is supported by the German Bundesministerium für Wirtschaft und Technologie (BMWi) via the Deutsches Zentrum für Luft- und Raumfahrt (DLR) under the contract numbers 50 QV 0301 and 50 OG 0502.

A.J.v.d.H. was supported by an appointment to the NASA Postdoctoral Program at the MSFC, administered by Oak Ridge Associated Universities through a contract with NASA. S.F. acknowledges the support of the Irish Research Council for Science, Engineering and Technology, cofunded by Marie Curie Actions under FP7. P.M. acknowledges the support of NASA NNX08AL40G. F.R. acknowledges the support of the Swedish National Space Board.

REFERENCES

- Abdo, A. A., et al. 2009, *ApJ*, 706, L138
 Ackermann, M., et al. 2010a, *ApJ*, 716, 1178
 Ackermann, M., et al. 2010b, *ApJ*, submitted
 Band, D., et al. 1993, *ApJ*, 413, 281
 Band, D. L., et al. 2009, *ApJ*, 701, 1673
 Bellm, E. C. 2010, *ApJ*, 714, 881
 Bhat, P. N. 2010, *GRB Circ.*, 10977, 1
 Crider, A., et al. 1997, *ApJ*, 479, L39
 Daigne, F., & Mochkovitch, R. 2002, *MNRAS*, 336, 1271
 Ford, L. A., et al. 1995, *ApJ*, 439, 307
 Gehrels, N. 1997, *Nuovo Cimento B*, 112, 11
 Ghirlanda, G., Celotti, A., & Ghisellini, G. 2003, *A&A*, 406, 879
 Golenetskii, S., et al. 2010, *GRB Circ.*, 10981, 1
 Goodman, J. 1986, *ApJ*, 308, L47
 Guiriec, S., et al. 2010, *ApJ*, 725, 225
 Kaneko, Y., Preece, R. D., Briggs, M. S., Paciesas, W. S., Meegan, C. A., & Band, D. L. 2006, *ApJS*, 166, 298
 Massaro, F., Grindlay, J. E., & Paggi, A. 2010, *ApJ*, 714, L299
 Meegan, C., et al. 2009, *ApJ*, 702, 791
 Mészáros, P. 2002, *ARA&A*, 40, 137
 Mészáros, P., & Rees, M. J. 2000, *ApJ*, 530, 292
 Pe'er, A., Ryde, F., Wijers, R. A. M. J., Mészáros, P., & Rees, M. J. 2007, *ApJ*, 664, L1
 Piran, T. 1999, *Phys. Rep.*, 314, 575
 Piran, T. 2004, *Rev. Mod. Phys.*, 76, 1143
 Preece, R. D., Briggs, M. S., Mallozzi, R. S., Pendleton, G. N., Paciesas, W. S., & Band, D. L. 1998, *ApJ*, 506, L23
 Preece, R. D., Briggs, M. S., Mallozzi, R. S., Pendleton, G. N., Paciesas, W. S., & Band, D. L. 2000, *ApJS*, 126, 19
 Rees, M. J., & Mészáros, P. 2005, *ApJ*, 628, 847
 Ryde, F. 2004, *ApJ*, 614, 827
 Ryde, F. 2005, *ApJ*, 625, L95
 Ryde, F., & Pe'er, A. 2009, *ApJ*, 702, 1211
 Ryde, F., et al. 2010, *ApJ*, 709, L172
 Tanaka, Y., Ohno, M., Takahashi, H., Uehara, T., Omodei, N., Chiang, J., & Guiriec, S. 2010, *GRB Circ.*, 10978, 1
 Woosley, S. E. 1993, *ApJ*, 405, 273
 Zhang, B., & Pe'er, A. 2009, *ApJ*, 700, L65